

JWST unveils the brown dwarf sequence of 47 Tucanae

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ABSTRACT

We have developed a technique to restore scientific usage in compromised (publicly available) images collected with the *JWST* of the Galactic globular cluster NGC 104 (47 Tucanae). In spite of the degradation and limited data, we were able to recover photometry and astrometry for the coolest stellar objects ever observed within a globular cluster, possibly unveiling the brightest part of the brown dwarf (BD) sequence. This is supported by (i) proper motion membership, derived by the comparison with positions obtained from *Hubble Space Telescope* archival early epochs, (ii) the predicted location of the BD sequence, and (iii) the mass function for low-mass stars derived from models. Future *JWST* observations will provide the necessary deep and precise proper motions to confirm the nature of the here-identified BD candidates belonging to this globular cluster.

Key words: brown dwarfs – globular clusters: individual: NGC 104.

1 INTRODUCTION

Globular clusters (GCs) have always been fundamental benchmarks to test models of stellar evolution, as they are formed by stars spanning a wide range of masses and with roughly the same age, chemical composition, and distance. GCs are among the oldest known objects of the Milky Way and possibly in the Universe, and their properties can be determined using stellar colour–magnitude diagrams (CMDs).

The *Hubble Space Telescope* (*HST*) has made significant contributions to our understanding of these systems, both in the Galactic (e.g. Bedin et al. 2004; Piotto et al. 2007) and extragalactic context (Nardiello et al. 2019). With its superior resolution and infrared (IR) sensitivity, we can expect that the *JWST* will also greatly enhance the study of GCs. Arguably, the most probable finding of *JWST* will be the complete characterization of GCs down to the entire lowest masses along the main sequence (MS), and in the so far unexplored regions of brown dwarfs (BDs).

BDs are objects that do not have enough mass to fuse hydrogen in their cores and instead contract to the size of Jupiter, cooling and dimming over time (Hayashi & Nakano 1963; Kumar 1963). While thousands of BDs of solar metallicity have been identified and studied near the Sun and in nearby young clusters and associations (Rebolo, Zapatero Osorio & Martín 1995; Kirkpatrick et al. 2019; Best et al. 2021), they have not yet been observed in GCs despite dedicated searches (Dieball et al. 2016, 2019). Identifying BDs in old and metal-poor GCs would greatly improve our knowledge of BDs and GCs. Indeed, the chemical composition of BDs has a major impact on their properties. For example, the mass at the hydrogen burning limit (HBL) is set by the internal opacity and metallicity: metal-free BDs have higher HBL mass ($0.090 M_{\odot}$) than BDs of solar abundance ($0.072 M_{\odot}$; Saumon et al. 1994). Cooling is also affected by the

chemical composition through opacity. Low temperatures allow for the formation of complex molecules in BDs' atmospheres, which in turn affects significantly their spectra (Birky et al. 2020). As BDs cool down (i.e. age), they increasingly separate in luminosity from stars. In the case of the extreme ages of GCs (> 10 Gyr), a significant gap in luminosity is expected, the width of which depends on the interior equation of state and low-temperature fusion processes. Studies of GCs with the *HST* have identified the beginning of this gap, but the most massive BDs (the brightest) at the bottom of the gap have not been detected yet. It is expected that the capabilities of *JWST* will reach these BDs in GCs, allowing for an independent measurement of the age of GC populations through the HBL gap (Caiazzo et al. 2017, 2019).

Concerning the multiple population phenomenon observed within GCs (Nardiello et al. 2015, 2018a; Piotto et al. 2015; Bastian et al. 2019), particularly in the most massive and dynamically less evolved ones (Bedin et al. 2004; Piotto et al. 2007; Siegel et al. 2007), BDs can provide important information. Indeed, BDs are highly sensitive to variations in composition and can serve as amplifiers of these differences, which can help to fine-tune models of the atmospheres of ultra-cool stars and substellar objects of different masses, ages, and chemical properties. This can help relate the properties of observed stellar populations at higher masses, in order to better understand the observed differences in chemical compositions among multiple generations of stars, which are not yet fully understood (Renzini et al. 2015).

In this letter, we derived photometry and astrometry for the coolest stellar objects ever observed in a GC, possibly unveiling for the first time the brightest part of the BD sequence in a GC.

2 OBSERVATIONS AND DATA REDUCTION

The data employed in this letter are *JWST* public archival material, which was compromised by technical problems. The programme is GO-2560 (PI: Marino), which focuses on characterizing multiple stellar populations in very low mass M dwarfs of the relatively close

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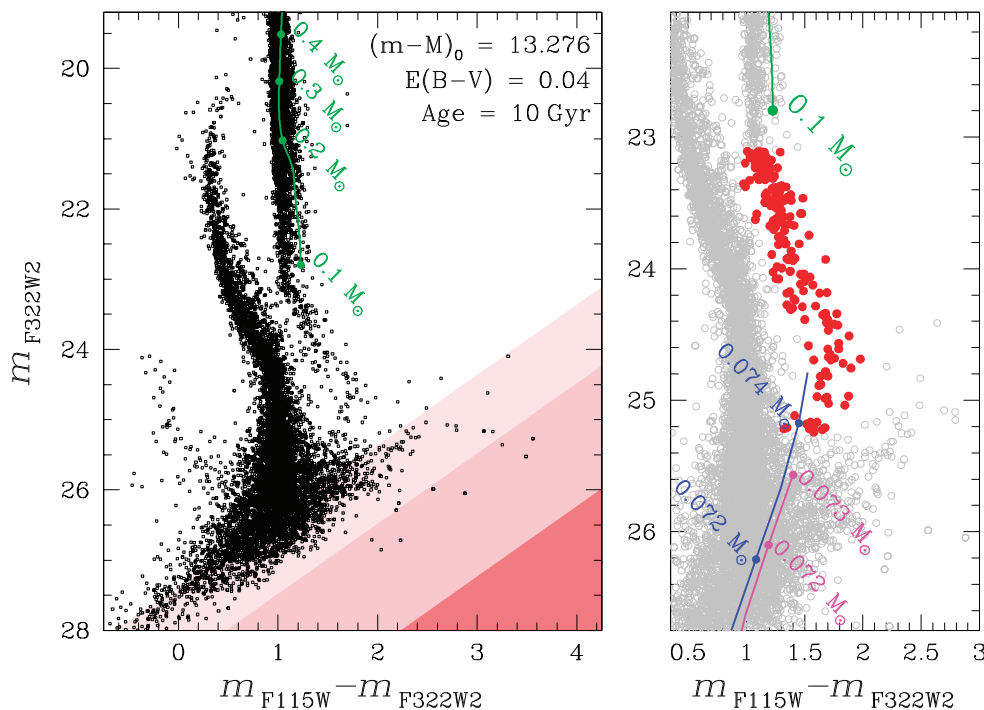


Figure 1. Left-hand panel: m_{F322W2} versus $(m_{F115W} - m_{F322W2})$ CMD for all the well-measured stars in the *JWST* field. Green line represents the 10 Gyr BaSTI-IAC isochrone. Shaded regions are our detection limits in both the filters: dark red corresponds to the background noise (1σ) converted in magnitude, and middle and light red are the 3σ and 5σ limits, respectively. Right-hand panel is a zoom of the CMD around the region where the high-mass BDs are expected. We highlighted in red the objects whose mass should be between 0.07 and $0.1 M_{\odot}$. In blue and magenta are reported the solar-metallicity BD 10 Gyr isochrones (ATMO2020 CEQ and NEQ strong models, respectively) from Phillips et al. (2020).

and massive Galactic GC 47 Tucanae (NGC 104). Visit 13 of GO-2560 suffered from a guide star acquisition failure, which has resulted in imperfect images. These images were judged unsuitable for the main science of that programme and requested (and approved) to be repeated in 2023 June. The discarded data, according to Space Telescope Science Institute (STScI) policies, become immediately public.

The data for this article were collected using the Near Infrared Camera (NIRCam) on *JWST* and centred on a region located about 6 arcmin from the 47 Tuc’s centre. The observations took place on 2022 July 13 and 14, and utilized two filters: F115W (short wavelength) and F322W2 (long wavelength); 40 exposures of 1030.7 s were taken with each filter, using the DEEP8 readout mode. The observations were collected with two large dithers to fill gaps between the detectors within a module, but not the gap between the two modules.

In our efforts to extract optimal astrometric and photometric capabilities of *JWST* cameras (Griggio, Nardiello & Bedin 2022; Nardiello et al. 2022), we searched the archive for dense stellar fields, and quickly realized that these failed images are an optimal benchmark to test our algorithms. Our procedures derive completely empirical effective point spread functions (ePSFs) of any shape, and rely on well-characterized geometric distortion for the cameras. In the following, we briefly describe the data reduction. First, taking advantage of the large dithers of the images and of the geometric distortion solution obtained by Griggio et al. (2022), we extracted a grid of 5×5 library ePSFs both for the F115W and F322W2 filters, following the procedure described by Nardiello et al. (2022). By using the brightest and most isolated stars in each image, we perturbed the ePSFs to take into account their time variations and we used them to obtain positions and fluxes of the sources detected in any given image (we refer to this as *first-pass* photometry). For each filter, we transformed positions and magnitudes in a common reference

frame, defined by the *Gaia* DR3 catalogue (Gaia Collaboration 2021) and the first image in each filter. We used images, perturbed ePSFs, and transformations to carry out the so-called *second-pass* photometry, by using the KS2 software, developed by J. Anderson (Anderson et al. 2008), which was employed and described in several works (e.g. Bellini et al. 2017; Nardiello et al. 2018b; Scalco et al. 2021), and here we have modified a private copy of the code to be adapted to NIRCam images. This routine, analysing all images simultaneously in a consistent reference frame, allows us to go deeper than the first-pass photometry and to detect extremely faint sources that would be otherwise lost in the noise of individual exposures. We carefully cleaned the final catalogue from the artefacts and bad sources (in part due to the imperfection of the images) by using the quality parameters output of the KS2 routine (Anderson et al. 2008). Calibration of the instrumental magnitudes were obtained as in Nardiello et al. (2022).

In this study, we made also use of the partially overlapped *HST* Wide Field Camera 3 (WFC3) optical and IR observations collected during the GO-11677 (PI: Richer, Kalirai et al. 2012) in F606W, F110W, and F160W. Astro-photometric catalogues were obtained from this data set following the data-reduction procedure described by Nardiello et al. (2018b) and Bedin & Fontanive (2018, 2020).

3 COLOUR-MAGNITUDE DIAGRAMS AND MOTIONS

In spite of the degradation suffered by these *JWST* images, our procedures were able to recover exquisite photometry and astrometry, which allowed us to undertake an independent investigation on a quite different subject of the nominal main goal proposed for the

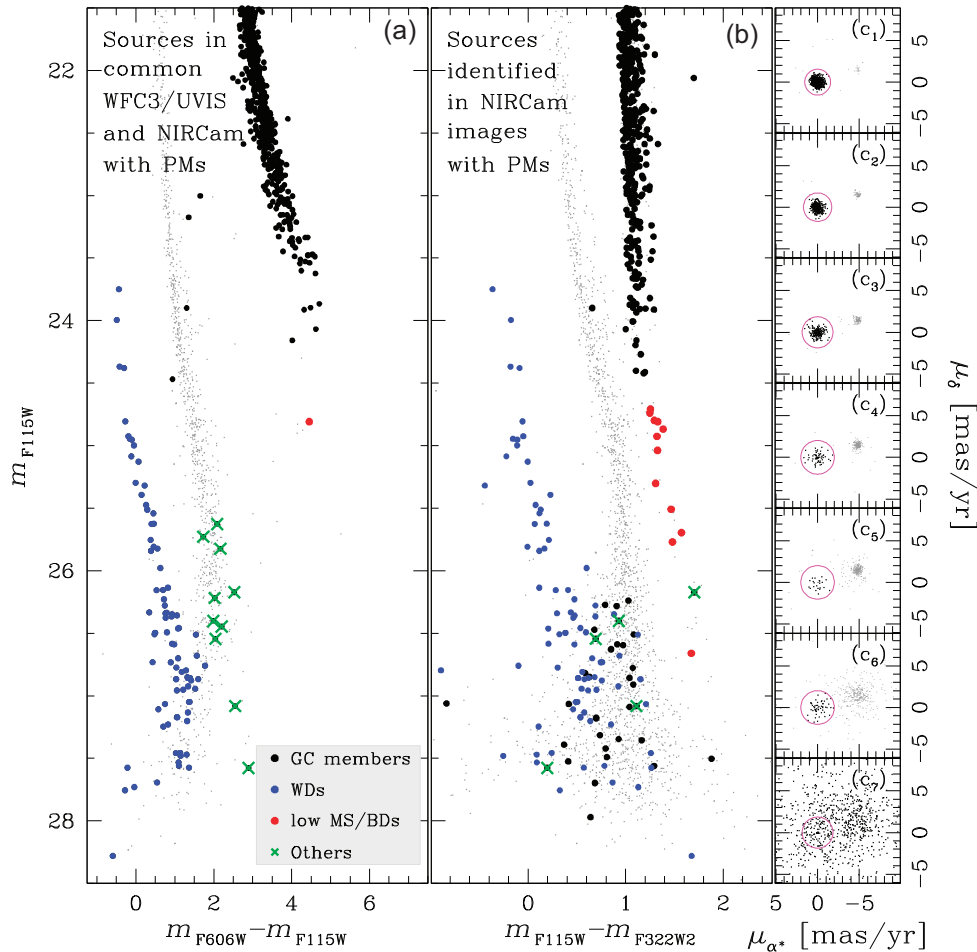


Figure 2. Identification of 47 Tuc’s members through PMs. Panel (a) shows the m_{F115W} versus $(m_{F606W} - m_{F115W})$ CMD, used to identify likely WDs (in blue) and faint objects visible in the F606W band (green crosses). These selections are used to exclude some candidate low MS/BDs (in red) in the m_{F115W} versus $(m_{F115W} - m_{F322W2})$ CMD of panel (b), which shows all the stars identified in the *JWST* data and for which we measured PMs with F606W, F110W, and F160W images. Panels (c) illustrate the PMs in the magnitude intervals corresponding to the ordinate axis of panels (a) and (b); black stars in the magenta circles are the likely GC members.

programme GO-2560. In this letter, we present our study of the BDs of 47 Tuc.

In left-hand panel of Fig. 1, we show the m_{F322W2} versus $(m_{F115W} - m_{F322W2})$ CMD for all stars in the field, where the cluster’s lower MS, the white dwarf (WD) sequence, and the Small Magellanic Cloud (SMC) MS are all well populated and clearly defined. Shaded regions show the 1σ , 3σ , and 5σ floor-noise level (computed as described in Bedin et al. 2023). We overlapped (in green) a BaSTI-IAC 10 Gyr isochrone (Hidalgo et al. 2018; Pietrinferni et al. 2021) to the MS of 47 Tuc. We used the distance modulus $(m - M)_0 = 13.276$ tabulated by Vasiliev & Baumgardt (2021) and the reddening $E(B - V) \sim 0.04$ as reported by Harris (1996). Beside usual uncertainties of models in reproducing the observed very low mass MS stars, it is also clear that the lower MS of 47 Tuc (down to $m_{F322W2} \sim 23$) is far from being well represented by a single stellar population (as expected). A more detailed analysis of the multiple population phenomenon in the lower MS of 47 Tuc will be the subject of a future article (Nardiello et al., in preparation). These theoretical models end at the lower limit $M = 0.1 M_{\odot}$ at magnitude $m_{F322W2} \sim 23$.

In right-hand panel of Fig. 1, we highlighted in red the stars that are likely cluster members and have masses $M < 0.1 M_{\odot}$. To interpret the substellar objects, we employed the BD tracks by

Phillips et al. (2020, ATMO2020) and computed 10 Gyr isochrones for solar-metallicity BDs (being not available to us BDs models at 47 Tuc metallicities). The line in blue shows the isochrone obtained from chemical equilibrium atmosphere (CEQ) models, while the one in magenta shows the one obtained employing atmosphere with non-equilibrium chemistry and ‘strong’ mixing (NEQ strong). Although these models do not reflect the exact chemistry of the stars in 47 Tuc, it is possible to note that the most massive BDs ($M > 0.07 M_{\odot}$) positions on the CMD are consistent with the observed sources highlighted in red.

In order to verify the cluster membership of these low-mass stars (indicated in red in Fig. 1), and the membership of even fainter stellar objects (which are completely lost in the MS of the SMC in these CMDs), we reduced deep *HST* observations collected with the WFC3 (see Section 2). These early epochs, along with the new *JWST* positions, were then employed to compute the relative proper motions (PMs) as in Bedin et al. (2014), by using 47 Tuc’s stars as reference system for the motion, and MS stars as reference for the local transformations. The average temporal baseline between the two epochs is $\Delta t = 12.4$ yr. The resulting PMs are reported in panels (c) of Fig. 2: we selected the likely cluster members by tagging the stars inside the magenta circles as sources with a PM in agreement

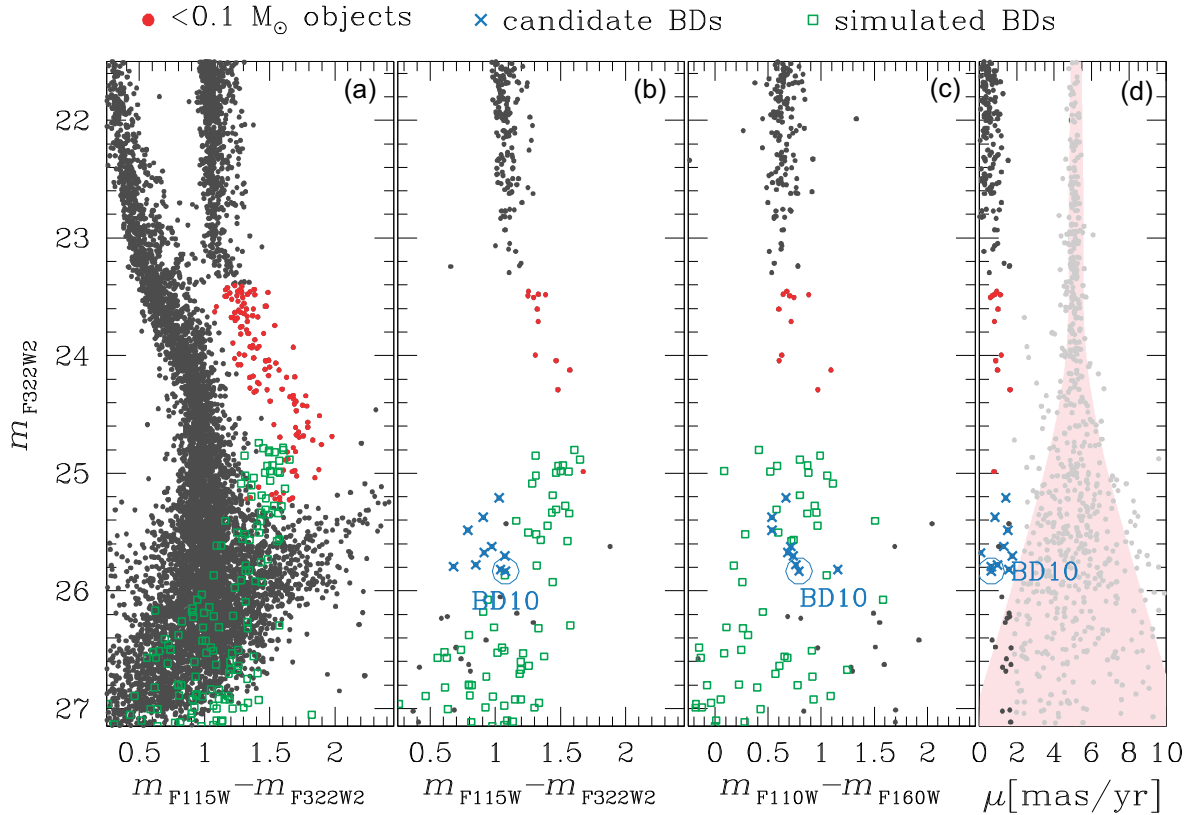


Figure 3. Comparison between simulated BDs (green squares) and observed CMDs. In panel (a), all the stars measured in the *JWST* images are plotted in black, while in red are indicated the objects consistent with mass $<0.1 M_{\odot}$. In next three panels, for clearness purposes, we do not plot the WDs identified in Fig. 2. Panels (b) and (c) show the m_{F322W2} versus $(m_{F115W} - m_{F322W2})$ and the m_{F322W2} versus $(m_{F110W} - m_{F160W})$ CMDs, respectively, for the stars with PMs in agreement with the mean cluster's motion. Azure crosses are the candidate BDs we identified in this study. Panel (d) shows the PM distribution referred to the mean cluster's motion. The shaded, rose region contains the stars whose PM is within 2σ from the mean motion of the SMC's stars. The azure circle is the faintest candidate BD we found, whose multifilter finding charts are shown in Fig. 4.

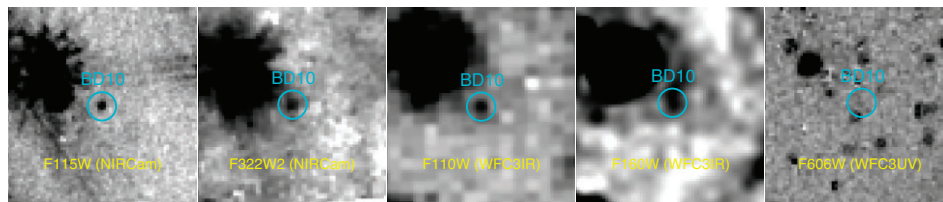


Figure 4. Finding charts of the faintest candidate BD found in this letter. The star is visible in all the adopted IR bands, but not in the optical band F606W, as expected for a BD.

with the cluster motion. Very red, cool objects are not visible in the m_{F115W} versus $(m_{F606W} - m_{F115W})$ CMD, and we used this CMD [panel (a) of Fig. 2] to identify, among the likely cluster members, the WDs (in blue) and other faint objects that cannot be low-MS stars/BDs (in green). In panel (b), we show the stars for which we have PMs; likely WDs and other objects are marked as in panel (a), while with a red circle we highlighted $<0.1 M_{\odot}$ stars and candidate BDs.

In the next section, we compare the observed CMDs with synthetic CMDs obtained from theoretical models.

4 COMPARISON WITH THEORETICAL ISOCHRONES

To obtain a qualitative estimate of the expected number of BDs at different magnitudes and their approximate location in the CMD, we first derived the mass function (MF) for the observed MS stars,

employing the BaSTI-IAC models. The derived masses turned out to be rather flat in the mass range $0.15\text{--}0.35 M_{\odot}$. We then extrapolated this derived flat MF value down to the BDs mass range ($0.015\text{--}0.075 M_{\odot}$), and computed a synthetic CMD (ignoring completeness) by employing the only BD isochrones at our disposal (described in previous section), which however are for solar metallicity. To generate the synthetic CMD, we proceeded as follows: we first generated a random $\log_{10}(\text{Mass})$ in the interval $[-1.8, -1.12]$ following a uniform distribution. For each random mass, we associated a magnitude in the *JWST* and *HST* filters using the BDs models, linearly interpolating between the two nearest theoretical points. We then added a random noise in magnitude with errors as estimated from the real data.

Synthetic BDs are shown with green symbols (squares) in Fig. 3, while the observed sources with black points. Panel (a) shows all the observed stars, while panels (b) and (c) only the subsample of

the observed stars with an estimate for the PMs, for the CMDs m_{F322W2} versus $m_{F115W} - m_{F322W2}$ (*JWST* only) and m_{F322W2} versus $m_{F110W} - m_{F160W}$ (*HST* only), respectively. Finally, panel (d) shows the combined 2D PMs as a function of *JWST*'s magnitude m_{F322W2} . This panel reveals an almost perfect separation between field objects (mainly SMC stars) and cluster members down to $m_{F322W2} \simeq 26$, below which the membership becomes less clear. We have identified (and highlighted with azure crosses) a group of 10 sources, which had PM consistent with the cluster's mean motion, survived to all the selections, and were located in the area of the CMDs where BDs with mass of $\sim 0.072 M_{\odot}$ are expected (synthetic BDs in green). We can also exclude that these objects are foreground BDs, as those would be close by, with considerably large dispersion and high PMs.

In summary, in this observational effort, we have employed compromised public images collected with NIRC*am* at the focus of *JWST*, recovering exquisite photometry and astrometry that enable scientific investigations. We used our reduced data to explore the faintest stellar objects in a field of 47 Tuc, where we isolated a group of 10 objects that we identify as candidate BD members of 47 Tuc. Their membership is supported by the following arguments: (i) their PMs are consistent with them being cluster members; (ii) their location on CMD is qualitatively consistent with the expected location by approximate models; and, finally, (iii) the observed number of candidate BDs is consistent with the simulated number obtained extrapolating the MS MF into the BD domain, also considering all the uncertainties (models, MF extrapolated, completeness, etc.).

Only adequate *JWST* follow-up observations will be able to confirm the nature of these BD candidates as true members of 47 Tuc, by means of multiwavelength deep observations and *JWST*-only derived PM.

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DATA AVAILABILITY

The data underlying this article are publicly available in the Mikulski Archive for Space Telescopes at <https://mast.stsci.edu/>. As electronic

material, we provide coordinates and magnitudes, for the 47 Tuc's BD candidates identified in this work.

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SUPPORTING INFORMATION

Supplementary data are available at *MNRASL* online.

47Tuc_BD_candidates.dat

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